

## The distant Type Ia supernova rate

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(The Supernova Cosmology Project)

### ABSTRACT

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We present a measurement of the rate of distant Type Ia supernovae derived using 4 large subsets of data from the Supernova Cosmology Project. Within this fiducial sample, which surveyed about 12 square degrees, thirty-eight supernovae were detected at redshifts 0.25–0.85. In a spatially-flat cosmological model consistent with the results obtained by the Supernova Cosmology Project, we derive a rest-frame Type Ia supernova rate at a mean redshift  $z \simeq 0.55$  of  $1.53^{+0.28}_{-0.25} {}^{+0.32}_{-0.31} 10^{-4} h^3 \text{Mpc}^{-3} \text{yr}^{-1}$  or  $0.58^{+0.10}_{-0.09} {}^{+0.10}_{-0.09} h^2 \text{SNU}$  (1 SNU = 1 supernova per century per  $10^{10} L_{B\odot}$ ), where the first uncertainty is statistical and the second includes systematic effects. The dependence of the rate on the assumed cosmological parameters is studied and the redshift dependence of the rate per unit comoving volume is contrasted with local estimates in the context of possible cosmic star formation histories and progenitor models.

## 1. Introduction

Recent observational efforts to detect high-redshift supernovae (SNe) have clearly demonstrated their value as cosmological probes. For the primary purpose of constraining the cosmic expansion history, the Supernova Cosmology Project (SCP) developed a scheduled search-and-follow-up technique that allows the systematic, on-demand discovery and follow up of “batches” of high-redshift SNe (Perlmutter et al. 1995a). Such batch discoveries of supernovae over the following years have led to the construction of two largely independent Hubble diagrams, one by the SCP (Perlmutter et al. 1997a, 1998, 1999) and one by the High-Z Supernovae Search Team (Garnavich et al. 1998; Schmidt et al. 1998; Riess et al. 1998), which both indicate significant, non-zero cosmological constant.

The batch discovery technique also provides well-controlled search conditions that make it possible to measure the *rate of occurrence* of distant SNe. In Pain et al (1996; hereafter Paper I), we presented the first such measurement using this technique. The distant supernova rate, and its comparison with the nearby supernova rate, can provide a diagnostic of the cosmic star formation history and metal enrichment at high-redshift, as well as a better understanding of possible SNe Ia progenitor models (Madau et al. 1998; Yungelson & Livio 1998). Obtaining a broader understanding of the nature and origin of high-redshift SNe will further improve and refine our use of supernovae as cosmological probes.

The local SNe Ia rate has recently been reported for two samples, one with  $z \simeq 0.01$  (Cappellaro et al. 1999) based on visual and photographic plates searches, and another at  $z \simeq 0.1$  (Hardin et al. 2000) based on CCD searches. In Paper I, we presented the SNIa rate at intermediate redshift ( $z$ ) ( $z \simeq 0.4$ ) using three SNe Ia discovered with the 2.5-m Isaac

Newton Telescope (INT). In this current paper, we report a refined measurement based on an enlarged sample of 38 SNe Ia, spanning the redshift interval 0.25–0.85, discovered over the course of four observing runs at the Cerro Tololo 4-m telescope. The new sample allows us, for the first time, to place constraints on the important question of possible evolution in the rate.

The method we adopt to calculate the SN rate is described in detail in Paper I, and contains two components. The first is the estimation of the SN detection efficiency and hence the “control time” (the effective time during which the survey is sensitive to a Type Ia event). We have studied our detection efficiency as a function of magnitude and supernova-to-host-galaxy surface brightness ratio using Monte-Carlo techniques. The second part estimates the comoving volume and total stellar luminosity to which our SNe survey is sensitive. We have computed the total galaxy luminosity from galaxy counts estimated from the Canada-France Redshift Survey (CFRS) and, independently, from recent parameterizations of the type-dependent field galaxy luminosity function and its redshift evolution. In combination, both aspects then yield an accurate determination of the SNIa rate at a mean redshift of  $z \simeq 0.55$ .

A plan of the paper follows. In §2 we discuss the new SN dataset and in §3 introduce our methodologies for estimating the control time and detection efficiencies. We reach significantly fainter detection limits compared to those of Paper I. In §4 we introduce the formalism for determining the survey comoving volume and in §5 various ways for estimating the accessible total stellar luminosity. This allows us to estimate the intermediate redshift SN rate in SNU (1 SNU = 1 supernova per century per  $10^{10} L_{B\odot}$ ). We discuss the various components of the uncertainties, statistical and systematic, in §6 and interpret our results in the context of local estimates and cosmic star formation histories in §7.

## 2. The Data Sets

For this analysis, we have studied 4 independent datasets of roughly equal size, totaling 219 similar search fields. These fields were observed in November and December of 1995 (Set A), in February and March of 1996 (Set B), in February and March of 1997 (Set C), and finally in November and December of 1997 (Set D), all using the Cerro Tololo 4-m telescope in Chile. The data sets were obtained as part of the search for high-redshift SNe conducted by the SCP. These images are suitable for a determination of the SN rate since they were obtained under similar conditions at one telescope, and therefore form well-defined, homogeneous sets.

Sets A and B were obtained using the 2048<sup>2</sup>-pixel prime-focus CCD camera, whereas Sets C and D were obtained with the 4 × 2048<sup>2</sup>-pixel Big Throughput Camera (BTC, Wittman et al. (1998)). The projected pixel size is  $\simeq 0.43''$  in both cases, giving an image size of approximately 16' × 16' (or 4 × 16' × 16' with the BTC). Exposure times were 2 × 600 s or more in the Kron-Cousins *R* filter, and the individual images reach a point-source 3 $\sigma$  magnitude limit ranging from  $R = 22.5$  mag to  $R = 24.5$  mag. Seeing was typically around 1''. The fields lie in the range  $0^h < \alpha < 15^h$ ,  $\delta > -10^\circ$ , avoiding the Galactic plane ( $|b| \gtrsim 30^\circ$ ). A few of the fields were selected due to the presence of a high-redshift cluster. The effect of the presence of clusters in the survey fields is taken into account in the calculation of the SN rate (see §4).

For all fields, a first-look “reference” image was obtained followed by a second look “search” image 2–3 weeks later. The useful solid angle of this dataset is defined by the overlap region of the original set of reference images with the search images. The total useful solid angle covered in this study is  $\simeq 12$  square degrees. The “reference” images were subtracted from the “search” images after convolution to match the seeing of the worst image and scaling in intensity. The resulting difference image for each field was searched for SN candidates. Tables 1a-d give the coordinates of the fields together with the supernova detection limit and the color excess  $E(B - V)$  derived from Schlegel, Finkbeiner & Davis (1998).

**Supernova Detection and Identification** The original search for supernovae was performed with a view to measure the cosmological parameters  $\Omega_M$  and  $\Omega_\Lambda$  (Perlmutter et al. 1999). The detection of supernovae was done in three steps: (1) the selection of transient events detected on the subtraction images with a signal-to-noise ratio cut of 3.5 $\sigma$  and a 15% cut on the ratio of the candidate flux and the host galaxy aperture flux at the candidate position. The latter cut had to be applied to remove systematics from subtraction residuals; (2) the rejection of statistical fluctuations, cosmic rays, asteroids with coincidences built from the multiple images of the same field taken at both epochs (“reference” and “search”); (3) the rejection of the remaining spurious candidates generated by hot or dead pixels, flatfield defects or bad subtractions with a visual inspection of each subtraction.

Altogether, 58 candidates passed the cuts in the original search and all but one were observed spectroscopically with the Keck Low-Resolution Imaging Spectrometer (LRIS, Oke et al. (1995)). The one remaining candidate was not followed up spectroscopically due to a lack of telescope time (and was thus not included in the cosmological parameter study in Perlmutter et al. (1999)). Its light curve, however, is consistent with that of a SNIa at redshift  $z \simeq 0.7$ . Of the 57 objects with spectral information, 4 were classified as “non

supernovae” (QSO/AGN) and the 53 remaining retained as possible supernovae (Perlmutter et al. 1995b, 1996, 1997b,c).

For the purpose of measuring the rate, a new search was performed on the same subtractions, slightly raising the signal-to-noise ratio cut (typically to  $5\sigma$ ) in order to ensure good control of the supernova detection efficiencies. Forty-six candidates remained at this stage (including the 4 “non supernovae”) of which 5 were spectroscopically identified as “non Ia” (II or QSO/AGN or Ib/c) and 37 as “possible SN Ia”. Type II supernovae were identified by the presence of hydrogen or by their very blue featureless spectrum, Ib/c by the absence of hydrogen and Si II or S II lines and the presence of narrow Ca II H&K features. The following criteria were then used to identify the SN Ia (Hook et al. 2002): (1) presence of Si II in the spectrum. For redshifts greater than  $z \sim 0.5$ , the Si II  $\lambda 4130\text{\AA}$  was used since Si II  $\lambda 6150\text{\AA}$  is beyond the spectroscopic range of *LRIS*; (2) presence of S II “W” feature at  $\sim 5500\text{\AA}$  when detected; (3) the large width of the  $\sim 4000\text{\AA}$  Ca II feature, characteristic of Type Ia SN.

Twenty eight candidates were identified as SN Ia using the above criteria leaving only nine for which the spectra had signal-to-noise ratio too low to distinguish among Type I sub-types. These 9 objects were discovered during the first 2 runs (Set A and set B) and observed spectroscopically under non optimal weather conditions. On the contrary, all objects discovered during the 2 other runs (Set C and set D) were observed with good signal-to-noise ratios. None of these events were classified as Ib/c. Considering the fact that all 4 sets have roughly equal sizes and were searched using the same procedures, this implies that the contamination by non-SN Ia in sets A and B is likely to be comparable, i.e. less than 10%. Two candidates have a E/S0 host type (Sullivan et al. 2002) which is a strong indicator of the supernova being of Type Ia. Adding the facts that the light curves of these partially identified objects resemble a Type Ia light curve at the observed redshift and that their peak magnitude is close to a Type Ia peak magnitude, we classified all 9 objects as “probable Ia”. These 9 events together with the one which was not observed spectroscopically were therefore retained for the rate analysis but the possibility that one of these objects may not be a SN Ia was used to estimate the effect of possible misidentification of supernova type on the systematic uncertainty (§6).

Altogether, thirty-eight SNe Ia with redshifts ranging from 0.25 to 0.85 were retained from the fifty-eight discovered. Redshifts were determined from spectra of the host galaxies. The properties of all 38 SNe Ia used in this analysis are summarized in Table 2.

### 3. Detection Efficiencies and Control Time

The data presented here were obtained with an observing strategy designed to measure the cosmological parameters  $\Omega_M$  and  $\Omega_\Lambda$  by conducting a search for supernovae on the rise using a subtraction technique. We followed the procedure introduced in Paper I to calculate the “control time” and detection efficiencies.

**Supernova Detection Efficiencies** Detection efficiencies were determined for every search field using Monte-Carlo simulations. A synthetic image was created for every field by adding simulated SNe to the search images. Reference images were subtracted from the synthetic search images using exactly the same software and cuts as used for the actual search, and the number of simulated SNe that satisfied the selection criteria was determined. The efficiency derived in this way then naturally accounts for parts of the image that are unusable for the SN search, for example regions saturated by bright foreground stars. Over two hundred simulated SNe were placed on each search image, with a range of apparent magnitude, host galaxy apparent magnitude and location with respect to host galaxies. Each simulated SN was generated by scaling down and shifting a bright star, with signal-to-noise ratio greater than 50, from the image being studied (it was not necessary to add additional Poisson noise because the dominant noise source is that of the sky). The position relative to the host galaxy was chosen at random from normal distributions with  $\sigma$  equal to the half width at half maximum of the galaxy independently on both  $x$  and  $y$  axis. The shift of the scaled bright star relative to the host galaxy was constrained to be an integral number of pixels in order to maintain the pixelized point spread function.

We reached significantly fainter detection limits during these observations compared to the data in Paper I. Figure 1 shows the fractional number of simulated SNe recovered, as a function of SN detected magnitude, for 12 representative examples among the 219 fields observed. For a typical field the detection efficiency is over 85% for any stellar object brighter than  $R = 23.5$ . Note that the loss in efficiency at the brightest magnitudes is due to detector saturation for bright sources. The plateau efficiency seen at intermediate magnitudes simply reflects the areal coverage lost due to masking of the region surrounding bright stars.

The efficiency depends primarily on the SN magnitude, but the Monte-Carlo simulation also permits to account for the small dependence of SN visibility on the host galaxy surface brightness underlying each SN. This is shown on Figure 2a where the overall supernova detection efficiency for Set A is plotted as a function of the magnitude difference between the host galaxy aperture flux at the supernova position and the supernova flux. Figure 2-b shows the overall supernova detection efficiency as a function of the projected distance to the host galaxy center. The detection efficiency does not depend on the SN position relative

to the host, demonstrating the ability of image subtraction techniques to detect supernovae on the nuclei of galaxies.

**Control Time** We computed a control time as a function of redshift and host galaxy magnitude equal to the weighted sum of the number of days during which the SN could be detected, given the time separation of the search and reference images, where the weighting is according to the corresponding detection efficiency.

Type Ia SN light curves are not unique. The total range for SN Ia  $B$ -band peak brightness spans  $\sim 0.5$  mag (Saha et al. 1999; Gibson et al. 2000). This has to be taken into account when computing the control time. Furthermore, as first noted by Phillips (1993), brighter supernovae also have wider light curves. This correlation between light curve shape and peak luminosity has the effect of further increasing the “visibility” of brighter objects and therefore the time during which they can be detected. To account for this correlation, the control time was computed, assuming that the SN Ia light curves form a one parameter family using an approximation for the light curve shape *vs.* luminosity relation following the “stretch factor” method of Perlmutter et al. (1997a). We assumed that the average SN light curve follows the average of the best-fit, time-dilated and  $K$ -corrected type Ia template (Leibundgut 1988), with the generalized cross-filter  $K$  correction described by (Kim, Goobar & Perlmutter 1996), and that the stretch parameter follows a Gaussian distribution with  $\sigma \sim 0.08$  (Perlmutter et al. 1999). The effect of the uncertainties in the light curve shape *vs.* luminosity correction and of the remaining  $\sim 0.15$  mag  $B$ -band peak luminosity intrinsic scatter on the systematic uncertainty in deriving the SN rates, is discussed in Section 6.

The SNe Ia light curves were calibrated using Landolt standards (Landolt 1992). Since these are observed light-curves, in *apparent* magnitudes, no explicit dependence of our rate on  $H_0$ ,  $\Omega_M$  or  $\Omega_\Lambda$  is introduced at this stage. Photometric calibration was not available for all the fields. For those fields without calibration (about 25%), zero points were calculated by comparison with  $E$ -band (which is close to  $R$ -band) magnitudes of anonymous stars in the APM (Automated Plate Measuring facility, Cambridge, UK) POSS-I catalog (McMahon & Irwin 1992). A comparison of the APM- $E$  magnitudes with CCD- $R$  magnitudes was performed using the fields on which SNe had been discovered. The distribution reveals a mean  $E - R$  offset of  $-0.02$  mag, with a dispersion of 0.22 mag. Assuming these fields are representative of the whole dataset, we applied a 0.02 mag shift to the APM magnitudes. The uncertainty in the rate introduced by this uncertain calibration is also discussed in Section 6.

Galactic extinction was taken into account for each field separately using two different methods. Firstly, we used the Galactic reddening value for each field  $E(B - V)$  supplied by

D. Burstein (private communication), derived from the analysis of Burstein & Heiles (1982). We applied these to the data assuming  $R_V = 3.1$ , and  $A_R/A_V = 0.751$  (Cardelli, Clayton & Mathis 1989). For the second method, we computed the extinction using more recent estimations of dust reddening (Schlegel, Finkbeiner & Davis 1998) and  $A_R/E(B - V) = 2.63$  (appropriate for the Landolt  $R$  filter). Although the individual field reddening values so determined can sometimes differ by a large amount even for our high latitude survey fields, the net effect on the rate is small as discussed in Section 6.

#### 4. SN Ia rates per unit comoving volume

To calculate the observed SN Ia rate per unit comoving volume, we derive the expected redshift distribution of SNe,  $N_{\text{exp}}(z)$ , which is proportional to the observed SN Ia rate,  $r_V(1+z)^{-1}$ , where  $r_V$  is the rest-frame SN rate per unit comoving volume and  $(1+z)^{-1}$  accounts for cosmological time dilation. The expected distribution is given by

$$N_{\text{exp}}(z) = \frac{r_V}{1+z} \sum_i S_i \times V(z; H_0, \Omega_M, \Omega_\Lambda) \times \Delta T_i(z) \quad (1)$$

where  $i$  runs over all the survey fields,  $S_i$  is the field solid angle, and  $V(z)$  is the comoving volume element at redshift  $z$  (formally  $d^2V/dz dS$ ) which depends on the cosmological parameters  $H_0$ ,  $\Omega_M$  and  $\Omega_\Lambda$  (see, for example, Equation (26) in Carroll, Press & Turner (1992)). Since the supernova detection efficiency depends on the galaxy apparent magnitudes ( $R_{gal}$ ), the control time per field at redshift  $z$ , ( $\Delta T_i(z)$ ) is computed as  $\Delta T_i(z) = \sum_R N_{gal}(z, R)_i \Delta T_i(z, R_{gal}) / \sum_R N_{gal}(z, R)_i$  where the sum runs over all possible galaxy apparent magnitudes. Individual control times have been calculated for each field in bins of  $z$  and  $R$  (the size of the bins used is 0.5 mag in  $R$  and 0.1 in  $z$ ).

Twenty seven of our 219 search fields had been chosen specifically to target high-redshift clusters. Suitable clusters and their redshifts were taken from (Gunn, Hoessel & Oke 1986). Although clusters will be found quite naturally in the SCP wide-field images, it is conceivable that they are over-represented. For each cluster target field, we therefore determined the excess number count as a function of  $R$  magnitude and cluster-centric radius by subtracting the appropriate background field. Assigning the known redshift of the appropriate cluster to the excess populations so determined, the effect on the rate per unit comoving volume was estimated by increasing the comoving volume element at the cluster redshift by the fractional excess of luminosity. The uncertainty in the rate introduced by targeting these high-redshift clusters is estimated in Section 6.



**One-parameter fits** Assuming a constant SN Ia rate as a function of redshift in the region covered by these data, we can perform a maximum likelihood fit of the observed redshift distribution to  $N_{\text{exp}}$  and hence derive  $r_V$  at a mean redshift,  $\bar{z} = \int z N_{\text{exp}}(z) dz / \int N_{\text{exp}}(z) dz$ .

The dependence of  $r_V$  on the Hubble parameter  $H_0$  is easily factorized (since the comoving volume element scales as  $H_0^{-3}$ ), but  $r_V$  also depends on the cosmological parameters  $\Omega_M$  and  $\Omega_\Lambda$ . At  $z = 0.5$ , the comoving volume element in a flat universe with  $\Omega_\Lambda = 0.7$  is twice that in a flat universe with no cosmological constant. Table 3 gives the results of the fits for different values of  $\Omega_M$  and  $\Omega_\Lambda$ . For a spatially-flat cosmological model with  $\Omega_M = 0.28$  as measured by the SCP (Perlmutter et al. 1999) and also reported by the High-Z search Team from their complete set of spectroscopic SNe Ia (Riess et al. 1998), we obtain

$$r_V = 1.53_{-0.25}^{+0.28}(\text{stat}) 10^{-4} h^3 \text{Mpc}^{-3} \text{yr}^{-1} \quad (2)$$

where the error is statistical only at this stage and  $h = H_0/100$ . Slightly different results for  $\Omega_M$  have also been reported by both groups depending on the sample retained in the analysis and the method used and  $\Omega_M$  is also been measured with different techniques (see for example Peacock et al. (2001)). It is therefore interesting to investigate the effect on the rate of changing the values of the cosmological parameters. A closer inspection of the comoving volume element dependence on  $\Omega_M$  and  $\Omega_\Lambda$  shows that, to a good approximation (< 5%), this quantity depends only on the difference  $\omega = \Omega_M - \Omega_\Lambda$  in our redshift range and for  $0.1 < \Omega_M < 1.5$  and  $|\omega| < 1.5$ . We therefore also provide the result as a function of  $h$  and  $\omega$  and find that the following is a good approximation to our results:

$$r_V(\bar{z} = 0.55) = [(2.06_{-0.33}^{+0.37}) \times (1 + 0.58\omega)] 10^{-4} h^3 \text{Mpc}^{-3} \text{yr}^{-1} \quad (3)$$

where the error is again only statistical at this stage.

A comparison of the expected number of SNe and the observed number is shown in Figure 3 where the expected number has been computed assuming no evolution for the rate per unit comoving volume and a flat universe with  $\Omega_M = 0.3$ . The agreement is quite good, although the expected distribution is slightly flatter.

Using the above determination of  $r_V$ , one can compute the theoretical number of SNe that are produced as a function of redshift. This is shown in Figure 4 where the number of SNe per square degree and per year is plotted as well as predictions for different cosmological models adjusted to best fit the observations (assuming that the number scales with comoving volume).

**Two-parameter fits** In the previous paragraph, the rest-frame SN rate  $r_V$  is assumed constant over the redshift range of interest. Several studies have addressed the expected

variation of the Ia rate with redshift (see for example Madau et al. (1998); Ruiz-Lapuente & Canal (1998); Sadat, Guiderdoni & Blanchard (1998)). With our enlarged sample spanning the redshift range 0.25–0.85, it is possible to consider an observational constraint on possible rate evolution. We choose to approximate any potential evolution with a power law of the form  $r_V(z) = r_{\bar{z}}[(1+z)/(1+\bar{z})]^\alpha$ , where  $r_{\bar{z}}$  is the  $z = \bar{z}$  type Ia rate per unit comoving volume and  $\alpha$  is an index of evolution ( $\alpha = 0$  indicates no evolution). Equation (1) then becomes

$$N_{\text{exp}}(z) = \frac{r_{\bar{z}}(1+z)^{\alpha-1}}{(1+\bar{z})^\alpha} \sum_i S_i \times V(z; H_0, \Omega_M, \Omega_\Lambda) \times \Delta T_i(z) \quad (4)$$

and we perform a two parameter fit of  $r_{\bar{z}}$  and  $\alpha$ .

As before, we perform maximum likelihood fits for a choice of cosmological models. The results are reported in Table 3. As expected, the evolution parameter  $\alpha$ , depends strongly on the assumed cosmology. A spatially-flat  $\Lambda$  dominated universe (model  $\Lambda$ ) favors a solution with little evolution in the SN Ia rate per unit comoving volume, whereas in an Einstein-de Sitter universe (model E) more evolution is permitted.

For the spatially-flat case (model  $\Lambda$ ) with  $\Omega_M = 0.28$ , we obtain (Fig. 5)

$$r_{\bar{z}=0.54} = 1.55^{+0.29}_{-0.30} 10^{-4} h^3 \text{ Mpc}^{-3} \text{ yr}^{-1} \quad \text{and} \quad \alpha = 0.8^{+1.6}_{-1.6} \quad (5)$$

where the error is statistical only.

Although the current dataset does not yet allow a stringent constraint to be placed on evolution in the SN Ia rate, with the ever-increasing number of SNe found in controlled experiments both at low and intermediate redshifts, the situation will improve quite rapidly. The SN Ia rate will therefore soon become a key ingredient in constraining the astrophysical evolution of host galaxies and in limiting possible progenitor models for SNe Ia.

## 5. SN Ia rates per unit galaxy luminosity

Local estimates of the SN Ia rate are often expressed in the “supernova unit” (SNU), the number of SNe per century per  $10^{10}$  solar luminosities in the rest-frame  $B$ -band. To compare our distant SN Ia rate with any local determinations, one must either convert the higher-redshift rates into SNU, or convert the local rates into “events  $\text{Mpc}^{-3} \text{ yr}^{-1}$ ”. In this section we explore the former option.

To estimate our rate in SNU, we proceed as described in Paper I and calculate the expected redshift distribution of SNe Ia given by

$$N_{\text{exp}}(z) = \frac{r_L}{1+z} \sum_i \sum_R N_{\text{gal}}(z, R)_i \times L_B(z, R; H_0, \Omega_M, \Omega_\Lambda) \times \Delta T_i(z, R) \quad (6)$$

where  $i$  runs over all fields,  $R$  is the galaxy apparent  $R$ -band magnitude, and  $L_B$  is the galaxy rest-frame  $B$ -band luminosity in units of  $10^{10} L_{B\odot}$  which depends upon the cosmological parameters  $H_0$ ,  $\Omega_M$  and  $\Omega_\Lambda$ .

Since thousands of anonymous high-redshift galaxies are observed in every survey image, it is more difficult than in local SNe searches to estimate the number, morphological type and luminosity distributions of galaxies searched within a given redshift range. To utilize a determination of the total  $B$ -band galaxy luminosity, as a function of  $z$  and apparent magnitude  $R$ , it will also be necessary to have the relevant galaxy  $K$ -corrections needed to convert observed  $R$  magnitudes into rest-frame  $B$  magnitudes.

We approach this determination of  $L_B(z, R)$  in two ways. Firstly, as in Paper I, we use observed  $R$ -band galaxy counts as a function of redshift and compute from these the rest-frame  $B$ -band galaxy luminosity. As a second estimate, we compute  $L_B(z, R)$  by integrating recently-determined luminosity functions parameterized via the Schechter function. We adopt  $M_{B\odot} = 5.48$ .

**Utilizing CFRS galaxy counts**  $R$ -band counts as a function of redshift were kindly calculated by Simon Lilly (Lilly 1995a) and are based on the analysis of  $I$ -band magnitude–redshift data obtained in the Canada-France Redshift Survey (Lilly et al. 1995b, and references therein). Since the  $I$ -band is fairly close to the  $R$ -band, and the magnitude range of the CFRS sample is comparable to that of our data, the extrapolation is small and therefore the dependence of  $R$ -band counts on the cosmological parameters is negligible. To compute the rest-frame  $B$ -band galaxy luminosities from apparent  $R$  magnitudes, we used  $B - R$  colors and  $B$ -band  $K$ -corrections provided by Gronwall (1995).

The SNIa rate per unit luminosity was then derived using this estimate of  $L_B(z, R)$  assuming that the rate per unit luminosity is constant as a function of redshift (an assumption we investigate in §6). The result is reported in Table 4.

**Utilizing observed luminosity functions** We also estimated  $L_B(z, R)$  by integrating recently-derived Schechter parameterizations of local field galaxy luminosity functions (LFs). We adopted a set of type-dependent LFs covering three broad galaxy classes; E/S0, spirals, and irregular systems. Many type-dependent LFs can be found in the literature (Loveday et al. 1992; Marzke et al. 1998; Folkes et al. 1999), based on many different surveys and classification techniques. The agreement is not particularly good and this renders our calculation somewhat uncertain (see, for example, Brinchmann (1999)). Bearing this in mind, we adopted the LFs of Marzke et al. (1998) as a reasonable “average”. Type-dependent  $K$

and luminosity evolutionary corrections were adopted from the synthesis models of Poggianti (1997). Finally, to apply these local LFs to higher-redshift samples, we also need to account for possible evolution in the LFs themselves. The primary signal is a marked increase with redshift in the abundance of galaxies with irregular morphology which we account for by introducing an evolution in the space-density of irregular systems, adjusted to match the evolution seen by Brinchmann et al. (1998).

Figure 6 shows the expected redshift distribution of SNe Ia,  $N_{\text{exp}}(z)$ , as calculated above for a spatially-flat  $\Lambda$  dominated cosmology (model  $\Lambda$ ) assuming that the SN Ia rate per unit luminosity does not evolve. Similar distributions were computed for different cosmological models and the rest frame SNe rate  $r_L$  was derived by fitting the redshift distribution of observed SNe to the expected distribution,  $N_{\text{exp}}(z)$ . Results of these fits are given in Table 4. For a flat universe with  $\Omega_M = 0.28$ , we find :

$$r_L(\bar{z} = 0.55) = 0.58_{-0.09}^{+0.10}(\text{stat}) h^2 \text{ SNU} \quad (7)$$

The value obtained for model  $\Lambda$  is in reasonable agreement (better than 10%) with the value obtained from the CFRS galaxy counts. This is because both estimates of the galaxy luminosity agree very well in the region  $z = 0.4$ – $0.6$ , where most of the SNe were found. Nevertheless, sizable differences exist in the high redshift region, where the luminosity derived from the CFRS counts lies significantly below that derived from the direct LF approach, probably due to the evolving population of (blue) irregular systems. Since a simple extrapolation was used to estimate the counts at high redshifts from the CFRS data, whereas the luminosity estimated from the parameterization of LFs used more recent high redshift survey data, the latter should be more realistic.

## 6. Systematic Uncertainties

With a total of almost forty SNe Ia, the statistical uncertainty is sufficiently small to demand a careful analysis of possible systematic uncertainties. We estimated these below and summarize their contribution in Table 5.

**Cosmological parameters** With the methods used in this paper to calculate the SN rates, the dependence on the cosmological parameters appears only in the calculation of the comoving volume element or in estimating the absolute galaxy luminosity. In both cases, the  $H_0$  dependence can be simply factorized. The dependence on  $\Omega_M$  and  $\Omega_\Lambda$  is more difficult to derive, although to a very good approximation ( $\simeq 5\%$ ) the comoving volume element depends

only on the combination  $\Omega_M - \Omega_\Lambda$  in our particular redshift range (assuming  $0.1 < \Omega_M < 1.5$  and  $|\Omega_M - \Omega_\Lambda| < 1.5$ , see §4). In the specific case of a spatially-flat cosmology, using the SCP value of  $\Omega_M = 0.28_{-0.09}^{+0.10}$  where statistical and systematic uncertainties have been combined, the uncertainty on the event rate becomes  ${}_{-0.23}^{+0.25} h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$ . For the rate per unit luminosity (i.e. in SNU), a simple parameterization on  $\Omega_M - \Omega_\Lambda$  is not possible, and for a flat universe with  $\Omega_M = 0.28_{-0.09}^{+0.10}$  we find a contribution of  ${}_{-0.03}^{+0.04} h^2 \text{ SNU}$ .

**Detection and identification efficiencies** The study of detection efficiencies as a function of SN magnitude is an essential element of this analysis. The detection efficiencies depend upon many parameters and vary widely from field to field. Uncertainties were determined using a statistically-limited Monte-Carlo simulation where 250 fake SNe were added to each image incorporating distribution functions for the galactocentric distance of each SN and the host galaxy magnitude distribution (assumed to be representative of the total galaxy population). The systematic uncertainties were estimated by varying the parameters of the simulation around their nominal values, and provide a fractional error on the efficiency of less than 5% or  $\pm 0.08 h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$  ( $\pm 0.03 h^2 \text{ SNU}$ ) for each contribution.

During the SCP SN search, differenced images of candidates satisfying loose cuts were scanned by eye and candidates kept or rejected following some quality criteria. This could give some systematic effects which are *a priori* difficult to estimate precisely. However for the rate analysis, an automatic procedure was used to retrieve the few hundred fake SNe that were added to each field in order to compute the detection efficiency. This makes it possible to estimate the SCP “scanning” efficiency. Interestingly, we found it to be better than 98% inside the nominal cuts which were set higher than during the actual search.

On the other hand, as discussed in Section 3, one SN candidate was not confirmed spectroscopically (see Table 2) and nine others were only spectroscopically identified as Type I and retained as “probable Ia” based on a combination of factors: two have E/S0 hosts and they all have light curve shape and magnitude at peak compatible with that of a Type Ia supernova. Furthermore, since no Ib/c was identified in the sets C and D from the 20 candidates which had spectra, the probability of having a Type Ib/c supernova in Set A and B, where the unidentified candidates have been found, is less than 10% with 90% confidence level. We therefore conclude that at most one of these 10 candidates which could not be positively identified as Ia could be a contaminant.

All together, we estimate the systematic uncertainty of both “scanning” efficiency and misclassification effects in the current dataset to be  $\pm 0.03 10^{-4} h^3 \text{ Mpc}^{-3} \text{ yr}^{-1}$  or  $\pm 0.01 h^2 \text{ SNU}$ . In combination, uncertainties in the detection efficiencies (detection, scanning, misclassification) translate into a overall systematic uncertainty on the rates of  $\pm 0.12$

$10^{-4} h^3 \text{ Mpc}^{-3} \text{ yr}^{-1} (\pm 0.04 h^2 \text{ SNU})$ .

**Range of SN Ia light curves** Control times were calculated following the procedure described in Section 3. Noticing that the different implementation of the light curve shape *vs.* luminosity correlation (Hamuy et al. 1996; Phillips et al. 1999; Riess, Press & Kirshner 1996; Riess et al. 1998; Perlmutter et al. 1997a) can give somewhat different corrections (Leibundgut 2001), we conservatively estimated the systematic uncertainty coming from the light curve shape *vs.* luminosity correlation by varying the “stretch” parameter by one standard deviation around its nominal value. The effect on the rate was found to be  $0.13 h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$  ( $0.04 h^2 \text{ SNU}$ ). The “intrinsic” scatter of 0.15 mag translates into a change in the rate of  $0.03 h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$  ( $0.02 h^2 \text{ SNU}$ ). The overall uncertainty due to the dispersion of SN Ia light curves therefore amounts to  $\pm 0.14 h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$  ( $\pm 0.05 h^2 \text{ SNU}$ ) on the rate.

**Field calibration** Measured SN light curves, calibrated with Landolt standards, were used to compute the control time. Efficiency versus magnitude curves, also needed to compute the control time, were obtained for all fields calibrated with Landolt standards when available (for 75% of the fields) or with the APM catalog for the others. Errors in the APM calibration of the fields thus alter the determination of the efficiency as a function of magnitude and therefore the control time. This has a sizable effect on the derived SN Ia rate since at the magnitude of most of our SNe, the detection efficiency varies rapidly with magnitude. We estimated the size of the effect by comparing the discovery magnitudes of our SNe calibrated using Landolt stars and calibrated with the APM, assuming that this was representative of our set of fields. Since only 25% of our fields lack Landolt calibration, the overall effect is reduced. It contributes  $\pm 0.06 10^{-4} h^3 \text{ Mpc}^{-3} \text{ yr}^{-1}$  ( $\pm 0.02 h^2 \text{ SNU}$ ) to the uncertainty in the rate.

**Galaxy luminosity** The CFRS galaxy counts are based on data that are well-matched to our survey in magnitude and redshift range, and only minimal extrapolation was required to convert from the *I* to *R* band. The associated uncertainty should be small and this is supported by the calculation based on using the observed LFs as discussed in Section 5 (Table 4). The difference in the two calculations serves as our estimate of the systematic uncertainty here and this amounts to  $\pm 0.05 h^2 \text{ SNU}$ .

**Contribution from clusters** Twenty seven of our 219 search fields had been chosen specifically to target high-redshift clusters. We followed the procedure described in Section 4

to account for the excess number counts that could arise from selecting these fields. Although the procedure may suffer large statistical and systematic uncertainties, it only affects a small fraction of the overall search area. We estimated a 50% overall uncertainty in estimating the excess number counts. This translates into less than a 10% uncertainty in calculating the overall contribution to the galaxy counts from clusters, giving a contribution of  $\pm 0.02 h^2 \text{ SNU}$  ( $\pm 0.05 10^{-4} h^3 \text{ Mpc}^{-3} \text{ yr}^{-1}$ ) to the uncertainty in the rate. It is likely this is an over-estimate of the uncertainty given we expect clusters to occur within typical survey fields.

**Galactic extinction** Galactic extinction was computed using two different methods, one taken from Burstein & Heiles (1982) based on emission from atomic neutral hydrogen, and the other from Schlegel, Finkbeiner & Davis (1998), based on dust emission in the far infrared (FIR). Both groups quote 10% uncertainty in their estimate of the reddening but differences as big as a factor of 2 in  $E(B - V)$  were found. However, the FIR emission maps have much better resolution which can be important in our case where the reddening has to be known for specific lines of sight. We therefore used Schlegel, Finkbeiner & Davis (1998) as our baseline. The effect on the rate is nevertheless very small since most of our fields were selected to have little or no reddening. Overall the  $\pm 10\%$  uncertainty in the reddening translates into an uncertainty on the rate of  $\pm 0.02 10^{-4} h^3 \text{ Mpc}^{-3} \text{ yr}^{-1}$  ( $\pm 0.01 h^2 \text{ SNU}$ ).

**Host galaxy inclination and extinction** The effect of host galaxy inclination on our detection efficiency and galaxy luminosity estimates should be taken into account when calculating SNe rates. Cappellaro et al. (1999) recently re-estimated the inclination correction factors for relevant nearby searches. In this analysis, both the search technique (in our case subtraction of CCD images) and calculation of the galaxies' luminosities were performed in a different manner than in most local searches, so the inclination effects may not be the same. Inclination and extinction would reduce both the number of SNe detected *and* the galaxy visible luminosity whose effects may partially cancel in estimating the rate. A complete analysis of this effect would require careful modeling of galaxy opacities, which is beyond the scope of this paper. Our result should therefore be directly compared with uncorrected values derived in nearby searches, with particular attention to CCD searches.

**Brightness evolution and intergalactic dust** The effect of possible SN Ia brightness evolution or presence of intergalactic dust was not explicitly taken into account in our derivation of the rates. However, since the SN Ia light curves used to compute the detection efficiencies were calibrated using the observed light curves, a possible difference in the brightness of distant SNe Ia compared to local ones is taken into account whether it is due to evolution

or cosmology. On the contrary, possible presence of intergalactic dust would have the effect of lowering the number of observed supernovae. In that case our results would have to be interpreted as a lower limit of the true distant rate.

## 7. Discussion

We have derived a rest-frame SNIa rate per unit comoving volume at redshift range 0.25–0.85 ( $\bar{z} \simeq 0.55$ ) of

$$r_V(\bar{z} = 0.55) = [(2.06_{-0.33}^{+0.37}(\text{stat}) \pm 0.20(\text{syst})) \times (1 + 0.58\omega)] 10^{-4} h^3 \text{Mpc}^{-3} \text{yr}^{-1} \quad (8)$$

with  $\omega = \Omega_M - \Omega_\Lambda$ , where the first uncertainty is statistical and the second includes systematic effects which are independent of the systematics arising from the uncertainty on the cosmological parameters.

For a spatially-flat universe consistent with the SCP results (i.e. with  $\Omega_M = 0.28_{-0.09}^{+0.10}$  (statistical and systematic uncertainty combined quadratically)), we measure :

$$r_V^{\text{flat}}(\bar{z} = 0.55) = 1.53_{-0.25}^{+0.28}(\text{stat})_{-0.31}^{+0.32}(\text{syst}) 10^{-4} h^3 \text{Mpc}^{-3} \text{yr}^{-1} \quad (9)$$

where the systematic uncertainty includes the uncertainty on the cosmological parameters.

As most low-redshift determinations of the SNIa rate are reported in ‘SNU’, we also estimate our SNe rate in these units. For the rate per unit luminosity, we obtain the following result

$$r_L(\bar{z} = 0.55) = 0.58_{-0.09}^{+0.10}(\text{stat})_{-0.09}^{+0.10}(\text{syst}) h^2 \text{SNU} \quad (10)$$

for a flat universe with  $\Omega_M = 0.28_{-0.09}^{+0.10}$ , and for an Einstein de Sitter universe, we measure  $r_L(\bar{z} = 0.55) = 0.94_{-0.14}^{+0.16}(\text{stat}) \pm 0.14(\text{syst}) h^2 \text{SNU}$  in good agreement with our first measurement reported in Paper I, based on the discovery of 3 SNeIa at  $z \simeq 0.4$ , of  $r_L(\bar{z} = 0.4) = 0.82_{-0.37}^{+0.54}(\text{stat})_{-0.25}^{+0.37}(\text{syst}) h^2 \text{SNU}$

We have studied the redshift dependence of the rate per unit comoving volume and put constraints on the rate of evolution of the SN Ia rate.

**Comparison with other estimates** In a recent work to be submitted for publication (Reiss 2002), D.J. Reiss reports values for the SN Ia rate per unit luminosity and per unit volume in excellent agreement with our values ( $< 1\sigma$ ). His values are based on a sample of 20 SNe at a mean redshift  $z \sim 0.49$ . Local  $z \simeq 0.01$  SN Ia rates have been recently reanalyzed, combining data from five SNe searches (see Cappellaro et al. (1999) and references therein).



They find  $r_L(z = 0.01) = 0.36 \pm 0.11 h^2 \text{ SNU}$ , averaged over all galaxy types. The quoted uncertainties include estimate of systematic effects. The SNIa rate at  $z \simeq 0.1$  has also been measured by Hardin et al. (2000), who find  $r_L(z = 0.1) = 0.44_{-0.21}^{+0.35} {}_{-0.07}^{+0.13} h^2 \text{ SNU}$  (here systematic and statistical errors are quoted).

In comparing these rates with our measurements, one should bear in mind the following caveats: (1) most local measurements (e.g. in Cappellaro et al. (1999) have been based on photographic data rather than CCD data as used here, (2) we did not apply any correction for host galaxy absorption and inclination, (3) at high redshift, the mix of galaxy types is likely to be very different (which will affect comparisons if different types have differing star-formation histories and hence SNIa rates), and (4) local SNIa rates are typically reported in SNU, whereas the high-redshift values are more conveniently calculated in “events  $\text{Mpc}^{-3} \text{ yr}^{-1}$ ” as the rest-frame  $B$ -band luminosity is difficult to estimate.

In this section, for the purpose of comparing to the models, we convert local rates from SNU to “events  $\text{Mpc}^{-3} \text{ yr}^{-1}$ ”. To do this, we calculate the  $B$ -band luminosity density of the local universe by integrating local  $B$ -band luminosity functions (Marzke et al. 1998; Folkes et al. 1999) and find  $\rho_{LB} = 1.7\text{--}2.7 \times 10^8 h L_{B\odot} \text{ Mpc}^{-3}$ ; we take an average value in this analysis, but note this introduces a further uncertainty into the calculation. We convert the local values and plot the results in Figure 7.

To this plot we have added recent theoretical predictions for the form that the evolution of the SNIa rate might take. Various workers have modeled the expected evolution of the SNIa rate (Ruiz-Lapuente, Canal & Burkert 1997; Sadat, Guiderdoni & Blanchard 1998; Madau et al. 1998; Sullivan et al. 2000a). However such work is hampered by the uncertain physical nature of the progenitor. The evolution expected depends critically on whether SNeIa occur in double or single degenerate progenitor systems (for a review see Nomoto et al. (1999) and references therein), the expected evolution in the fraction of stellar binaries and, of course, the cosmic star formation history (SFH).

Here we adopt an empirical approach representing these uncertainties in terms of two parameters (Madau et al. 1998; Dahlen & Fransson 1999). The first is a delay time  $\tau$ , between the binary system formation and SNe explosion epochs, which defines a (time-independent) explosion probability per white dwarf. Note that this parameter is treated in different ways in the literature. Madau et al. (1998) define a (time-independent) explosion probability per white dwarf, which they model as an exponential probability function with a mean value of  $\tau$  whereas Dahlen & Fransson (1999) use  $\tau$  as an *exact* delay time between binary system formation and supernovae explosion. The difference between the two approaches becomes sizable at higher redshift for scenarios involving large values of  $\tau$ . Here we adopt the former approach but note that this may introduce further uncertainties at  $z > 1$  in scenarios with

large  $\tau$ . The second parameter is an explosion efficiency,  $\eta$ , which accounts for the fraction of binary systems that never result in a SNe. We constrain  $\eta$  by requiring that our predicted rate at  $z = 0.55$  is equal to our new observational determination. In Figure 6, we show two illustrative values:  $\tau = 0.3$  Gyr, corresponding to a shallower decline at high-redshift, and  $\tau = 3.0$  Gyr which produces a steeper drop-off.

We consider each of these two SNe Ia models in the context of two different star formation history scenarios. The first (SFH-I) is taken from Madau & Pozzetti (2000) who provides a convenient analytical fit of the star formation rate (*SFR*) form

$$SFR(z) = \frac{1.67 \times 0.23 e^{3.4z}}{e^{3.8z} + 44.7} M_{\odot} \text{ yr}^{-1} \text{ Mpc}^{-3} \quad (11)$$

in an Einstein-de Sitter universe. We converted this formula to that appropriate for a  $\Lambda$ -dominated flat universe by computing the difference in luminosity density. The SFH fit matches most *UV*-continuum and  $H\alpha$  luminosity densities from  $z = 0$  to  $z = 4$  and includes a mild correction for dust of  $A_{1500} = 1.2$  mag ( $A_{2800} = 0.55$  mag). However, the *SFR* evolution in this model to  $z \simeq 1.5$  is both stronger, and results in a lower local *SFR*, than some recent *UV* measurements (Cowie, Songaila & Barger 1999; Sullivan et al. 2000b). Accordingly, we also consider a second SFH (SFH-II) with a shallower evolution (a factor of  $\sim 4$  from  $z = 0$  to  $z \simeq 1.75$  in an Einstein-de Sitter universe, and constant thereafter).

These various predictions are plotted, for a flat  $\Omega_M = 0.3$  cosmology, in Figure 7, together with our estimate of the evolutionary index. Although our internal estimate is highly uncertain, already it would seem to favor scenarios which involve little evolution over the redshift range  $z = 0$ – $0.6$ ; a result which is in agreement with comparisons based on the low redshift rate determinations. Clearly, a precise measurement of the SN Ia rate at (say)  $z = 1$  would enable further, more robust, constraints to be placed on any evolution, as the redshift range that is currently probed is quite small. In the near future, our Supernova Cosmology Project’s ongoing high redshift supernova searches, and those of the High-Z Search Team, should provide enough data at these redshifts to place more stringent constraints on the star formation history.

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Table 1a. Data Set A: fields A-1 to A-23

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
A-1	1 04 18.51	7 46 03.9	22.1	0.025
A-2	3 07 51.40	10 39 42.9	22.7	0.099
A-3	3 36 59.00	0 25 12.7	22.8	0.114
A-4	3 15 48.75	−1 34 39.4	23.1	0.031
A-5	3 42 32.38	17 30 38.9	22.5	0.128
A-6	1 56 55.65	7 42 58.2	23.1	0.040
A-7	1 55 00.61	7 53 37.1	23.0	0.046
A-8	1 56 28.73	8 07 11.5	23.1	0.043
A-9	1 54 47.22	7 55 15.0	23.0	0.041
A-10	1 54 38.32	7 59 21.6	23.2	0.044
A-11	1 54 28.74	8 17 10.7	23.2	0.044
A-12	1 55 45.31	8 15 19.3	23.1	0.038
A-13	1 53 21.74	7 33 23.9	23.0	0.033
A-14	1 53 32.34	7 57 34.0	22.9	0.038
A-15	1 53 34.85	8 18 02.8	23.0	0.052
A-16	1 33 57.62	6 20 25.0	23.5	0.044
A-17	1 43 10.43	2 32 17.6	23.0	0.023
A-18	1 49 38.77	2 04 49.0	22.9	0.048
A-19	2 03 05.76	1 55 26.9	23.0	0.029
A-20	2 06 59.11	6 52 13.1	22.7	0.035
A-21	2 37 45.13	3 43 23.6	22.6	0.035
A-22	4 59 02.18	7 57 55.9	23.2	0.235
A-23	5 16 46.29	14 48 35.0	23.4	0.688

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).

Table 1a. Data Set A: fields A-24 to A-46

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
A-24	8 55 10.82	8 01 16.7	23.2	0.018
A-25	8 26 59.23	4 35 37.6	23.9	0.027
A-26	8 52 05.41	2 15 22.2	21.8	0.034
A-27	2 08 11.00	−13 29 16.8	23.2	0.019
A-28	1 35 40.23	4 23 32.5	23.0	0.023
A-29	1 37 05.37	4 17 45.2	23.2	0.022
A-30	1 37 37.86	4 19 14.0	23.1	0.021
A-31	1 38 50.55	4 21 07.5	23.2	0.019
A-32	1 39 59.01	4 21 19.0	23.1	0.021
A-33	1 40 51.70	4 21 45.3	23.2	0.022
A-34	1 35 43.81	4 30 38.4	22.9	0.025
A-35	1 36 28.77	4 33 17.2	22.9	0.021
A-36	1 37 24.50	4 36 01.0	23.1	0.021
A-37	1 38 41.33	4 27 02.2	22.4	0.019
A-38	1 39 16.44	4 37 20.6	23.2	0.022
A-39	1 40 44.93	4 32 54.6	23.2	0.019
A-40	3 00 56.51	0 28 36.3	22.9	0.041
A-41	3 03 00.71	0 36 25.2	22.7	0.031
A-42	3 03 58.40	0 29 33.5	23.1	0.029
A-43	3 00 27.57	0 52 40.2	23.1	0.041
A-44	3 01 59.24	0 51 06.8	23.0	0.033
A-45	3 02 36.56	0 49 50.4	22.8	0.033
A-46	3 01 43.69	1 01 21.7	22.9	0.036

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).

Table 1a. Data Set A : fields A-47 to A-69

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
A-47	3 22 18.03	−4 58 15.8	23.1	0.032
A-48	3 23 13.26	−4 58 00.5	23.1	0.034
A-49	5 12 33.79	−5 28 24.2	23.5	0.094
A-50	5 14 08.04	−5 25 26.9	23.0	0.146
A-51	5 15 06.22	−5 22 47.4	23.5	0.142
A-52	5 15 42.22	−5 27 40.8	23.3	0.174
A-53	5 16 38.77	−5 19 49.7	23.5	0.184
A-54	5 17 29.27	−5 23 23.6	23.3	0.191
A-55	5 18 32.34	−5 27 45.2	23.1	0.199
A-56	5 11 55.21	−5 08 31.1	23.5	0.087
A-57	5 13 13.89	−5 15 08.8	23.5	0.125
A-58	5 14 28.75	−5 15 52.9	23.4	0.155
A-59	5 15 56.36	−5 09 13.4	23.5	0.132
A-60	5 16 21.30	−5 07 41.7	23.5	0.120
A-61	5 15 19.46	−4 52 38.5	23.4	0.123
A-62	5 15 26.23	−4 58 06.4	23.2	0.117
A-63	5 16 54.53	−4 55 53.2	23.5	0.091
A-64	8 13 58.35	10 02 08.8	23.2	0.038
A-65	8 16 03.00	10 02 51.0	23.3	0.040
A-66	8 17 32.60	10 07 47.0	23.0	0.035
A-67	8 14 56.29	10 11 05.5	22.7	0.042
A-68	8 15 42.78	10 22 30.0	23.2	0.039
A-69	8 16 58.45	10 45 50.5	23.2	0.037

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).

Table 1b. Data Set B: fields B-1 to B-23

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
B-1	12 40 43.23	−7 09 48.5	23.4	0.038
B-2	12 34 43.21	−9 24 52.4	23.3	0.036
B-3	11 21 33.31	0 07 09.2	23.5	0.043
B-4	10 40 17.51	−6 59 30.4	22.8	0.049
B-5	08 54 58.96	8 09 17.1	23.0	0.017
B-6	10 16 42.40	−1 10 36.9	22.4	0.031
B-7	08 51 34.90	2 16 35.6	22.9	0.032
B-8	09 00 20.78	3 53 52.6	22.6	0.036
B-9	12 26 48.93	11 16 46.7	22.6	0.033
B-10	12 57 57.65	−0 38 19.8	23.2	0.029
B-11	11 32 24.04	−3 07 30.4	23.3	0.035
B-12	13 17 29.17	−4 16 06.3	22.3	0.024
B-13	14 18 44.14	2 52 33.1	23.1	0.027
B-14	14 19 32.93	2 59 42.6	23.1	0.026
B-15	14 21 00.04	2 53 38.3	23.1	0.024
B-16	14 21 19.88	2 55 10.4	23.1	0.026
B-17	14 22 58.40	2 58 46.6	23.0	0.027
B-18	14 23 54.41	2 57 59.6	23.1	0.026
B-19	14 24 06.19	2 57 29.6	23.4	0.026
B-20	15 04 35.18	2 55 42.4	23.3	0.028
B-21	15 05 51.60	2 53 51.3	23.4	0.030
B-22	15 06 13.93	2 56 24.9	23.3	0.029
B-23	09 56 32.84	3 16 54.2	22.6	0.037

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).

Table 1b. Data Set B: fields B-24 to B-46

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
B-24	9 57 24.31	3 20 11.9	22.2	0.042
B-25	9 58 19.95	3 20 54.6	22.8	0.043
B-26	9 56 44.60	3 08 34.8	22.8	0.038
B-27	10 31 46.59	0 06 42.0	22.6	0.038
B-28	10 30 51.46	−0 06 44.0	22.5	0.041
B-29	11 23 37.67	0 47 12.5	23.4	0.041
B-30	11 24 39.81	0 43 27.8	23.1	0.040
B-31	13 17 50.56	−0 09 31.9	23.0	0.021
B-32	13 19 39.82	−0 06 45.5	22.6	0.024
B-33	13 19 59.64	−0 07 03.1	23.2	0.025
B-34	13 21 22.75	−0 08 11.1	23.4	0.025
B-35	13 22 20.39	−0 07 01.9	23.6	0.027
B-36	13 23 04.56	−0 07 10.9	23.5	0.033
B-37	13 24 26.99	−0 06 39.7	23.6	0.034
B-38	16 06 06.38	6 40 14.7	22.9	0.045
B-39	16 05 59.89	6 23 30.0	23.3	0.049
B-40	16 07 16.97	6 26 15.2	22.9	0.046
B-41	16 08 38.64	6 29 30.6	23.0	0.052
B-42	16 09 07.04	6 22 04.6	23.2	0.051
B-43	16 09 43.78	6 26 41.9	22.9	0.051
B-44	16 10 22.41	6 01 20.1	23.0	0.048
B-45	16 10 47.22	5 58 39.0	23.4	0.050
B-46	16 11 59.78	6 00 36.4	23.1	0.060

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).



Table 1c. Data Set C

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
C-1 <sup>b</sup>	8 15 49.75	10 00 22.4	23.8	0.040
C-5	8 56 15.99	4 41 47.7	23.1	0.021
C-9	8 59 04.49	4 39 53.8	23.4	0.027
C-13	8 58 34.19	4 00 32.8	22.9	0.029
C-17	11 23 28.58	0 56 39.2	24.2	0.043
C-21	11 31 30.08	−2 45 35.0	24.0	0.041
C-25	11 33 28.88	−2 42 35.2	24.0	0.037
C-29	11 31 22.77	−3 17 59.8	24.1	0.034
C-33	13 20 22.31	0 01 09.1	24.1	0.025
C-37	13 22 37.16	0 03 11.9	24.5	0.026
C-41	14 22 02.21	2 51 51.7	24.0	0.026
C-45	14 24 46.62	2 55 49.2	23.9	0.024
C-49	8 29 48.58	5 00 52.2	19.6	0.019
C-53	10 32 16.36	−0 12 47.3	23.2	0.045
C-57	10 35 10.67	0 27 23.7	23.3	0.031

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

<sup>b</sup> First CCD of the 4 2k×2k Big Throughput Camera.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).

Table 1d. Data Set D

Name	RA(2000)	DEC(2000)	Detection Limit <sup>a</sup>	E(B-V)
D-1 <sup>b</sup>	8 58 47.18	4 27 12.9	24.3	0.026
D-5	9 01 26.27	4 27 37.9	24.5	0.027
D-9	9 01 41.03	3 49 21.3	24.6	0.048
D-13	5 37 35.18	−2 53 03.5	22.7	0.048
D-17	5 37 33.80	−3 30 45.1	18.6	0.065
D-21	5 35 40.87	−2 26 18.1	24.0	0.058
D-25	5 35 37.78	−2 57 02.2	24.1	0.051
D-29	5 34 46.79	−3 27 48.5	24.4	0.045
D-33	5 33 31.54	−2 14 40.9	24.1	0.056
D-37	8 59 20.58	3 55 52.7	24.2	0.031
D-41	8 57 00.27	4 01 24.3	23.6	0.026

<sup>a</sup> Defined as the magnitude above which the supernova detection efficiency drops below 50% of the maximum detection efficiency in the field.

<sup>b</sup> First CCD of the 4 2k×2k Big Throughput Camera.

Note. — Col 1: Field Name, Col 2: Right Ascension (equinox 2000), Col 3: Declination (equinox 2000), Col 4: Supernova detection limit, Col 5: Color excess from Schlegel, Finkbeiner & Davis (1998).

Table 2. Thirty eight distant Type Ia Supernovae

Name	Redshift	Discovery <i>R</i> mag.	Name	Redshift	Discovery <i>R</i> mag.
1995aq	0.453	22.4	1997ag	0.592	23.2
1995ar	0.497	23.1	1997ai	0.450	22.3
1995as	0.498	23.3	1997aj	0.581	23.8
1995at	0.655	22.7	1997ak	0.347	24.4
1995aw	0.400	22.5	1997al	0.621	23.8
1995ax	0.615	22.6	1997am	0.416	23.4
1995ay	0.480	22.7	1997ap	0.830	24.2
1995az	0.450	24.0	unnamed <sup>a</sup>	~0.7	23.5
1995ba	0.388	22.6	1997el	0.636	23.1
1996cf	0.570	22.7	1997em	0.460	23.6
1996cg	0.460	22.1	1997ep	0.462	22.4
1996ch	0.580	23.7	1997eq	0.538	22.4
1996ci	0.495	22.3	1997er	0.466	22.3
1996ck	0.656	23.5	1997et	0.633	23.4
1996cl	0.828	23.6	1997eu	0.592	22.4
1996cm	0.450	22.7	1997ex	0.361	21.4
1996cn	0.430	22.6	1997ey	0.575	22.9
1997ac	0.320	23.1	1997ez	0.778	23.4
1997af	0.579	23.7	1997fa	0.498	22.5

<sup>a</sup>Not observed spectroscopically (see text)

Note. — Cols 1 & 5: IAU Name, Cols 2 & 5: Geocentric redshift of supernova or host galaxy, Cols 3 & 6: Approximate discovery *R* magnitude

Table 3. SN Ia rate per unit comoving volume for different cosmological models: in a flat  $\Lambda$  dominated model consistent with the latest Supernova Cosmology Project results (Model  $\Lambda$ ); in a  $\Lambda = 0$  universe with  $\Omega_M = 0.3$  (Model O); and in an Einstein-de-Sitter universe (Model E).

Model	$\Omega_M$	$\Omega_\Lambda$	$\bar{z}_{\text{exp}}^{\text{a}}$	$\bar{z}_{\text{obs}}^{\text{b}}$	$r_V^{\text{c}}$	$\alpha^{\text{d}}$
<i>One parameter Fits</i>						
$\Lambda$	0.28	0.72	0.53	0.54	$1.53_{-0.25}^{+0.28}$	...
O	0.3	0.0	0.52	0.54	$2.42_{-0.40}^{+0.44}$	...
E	1.0	0.0	0.52	0.54	$3.25_{-0.53}^{+0.58}$	...
<i>Two parameters Fits</i>						
$\Lambda$	0.28	0.72	0.54	0.54	$1.55_{-0.30}^{+0.29}$	$0.8_{-1.6}^{+1.6}$
O	0.3	0.0	0.54	0.54	$2.48_{-0.48}^{+0.48}$	$1.3_{-1.6}^{+1.6}$
E	1.0	0.0	0.54	0.54	$3.36_{-0.64}^{+0.64}$	$1.7_{-1.6}^{+1.5}$

<sup>a</sup>Expected mean redshift. Computed from the expected number of supernovae  $N_{\text{exp}}(z)$  (see text).

<sup>b</sup>Observed mean redshift

<sup>c</sup>Rate per unit volume ( $h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$ ) at mean redshift  $z = \bar{z}_{\text{exp}}$ . Statistical uncertainty only

<sup>e</sup>Evolution index (see text and Figure 4)

Table 4. SN Ia rate per unit luminosity for different cosmological models: in a flat  $\Lambda$  dominated model consistent with the latest Supernova Cosmology Project results (Model  $\Lambda$ ); in a  $\Lambda = 0$  universe with  $\Omega_M = 0.3$  (Model O); and in an Einstein-de-Sitter universe (Model E).

Model	$\Omega_M$	$\Omega_\Lambda$	$\bar{z}_{\text{exp}}^{\text{a}}$	$\bar{z}_{\text{obs}}^{\text{b}}$	$r_L^{\text{a}}$
<i>From CFRS galaxy counts</i>					
$\Lambda$	0.28	0.72	0.56	0.54	$0.63_{-0.10}^{+0.11}$
<i>From Luminosity Functions</i>					
$\Lambda$	0.28	0.72	0.58	0.54	$0.58_{-0.09}^{+0.10}$
O	0.3	0.0	0.57	0.54	$0.78_{-0.13}^{+0.14}$
E	1.0	0.0	0.57	0.54	$0.91_{-0.14}^{+0.16}$

<sup>a</sup>Expected mean redshift. Computed from the expected number of supernovae  $N_{\text{exp}}(z)$  (see text).

<sup>b</sup>Observed mean redshift

<sup>c</sup>Rate per unit luminosity ( $h^2$  SNU) at mean redshift  $z = z_{\text{exp}}$ . Statistical uncertainty only

Table 5. Summary of Uncertainties

Source	$\delta r_V^a$	$\delta r_L^b$
Cosmological parameters	+0.25 -0.23	+0.04 -0.03
Detection efficiencies	$\pm 0.12$	$\pm 0.04$
Range of Ia lightcurves	$\pm 0.14$	$\pm 0.05$
Field calibration	$\pm 0.06$	$\pm 0.02$
Cluster contribution	$\pm 0.05$	$\pm 0.02$
Galaxy extinction	$\pm 0.02$	$\pm 0.01$
Luminosity estimate	$\dots$	$\pm 0.05$
Total syst. uncertainty	+0.32 -0.31	+0.10 -0.09
Statistical uncertainty	+0.28 -0.25	+0.10 -0.09

<sup>a</sup>Uncertainty on the rate per unit volume ( $h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$ )

<sup>b</sup>Uncertainty on the rate per unit luminosity ( $h^2 \text{ SNu}$ )

Note. — These uncertainties have been computed in a flat  $\Lambda$  dominated universe using  $\Omega_M = 0.28_{-0.09}^{+0.10}$  (see text). No estimate was made of possible systematic uncertainties from host galaxy inclination or extinction.

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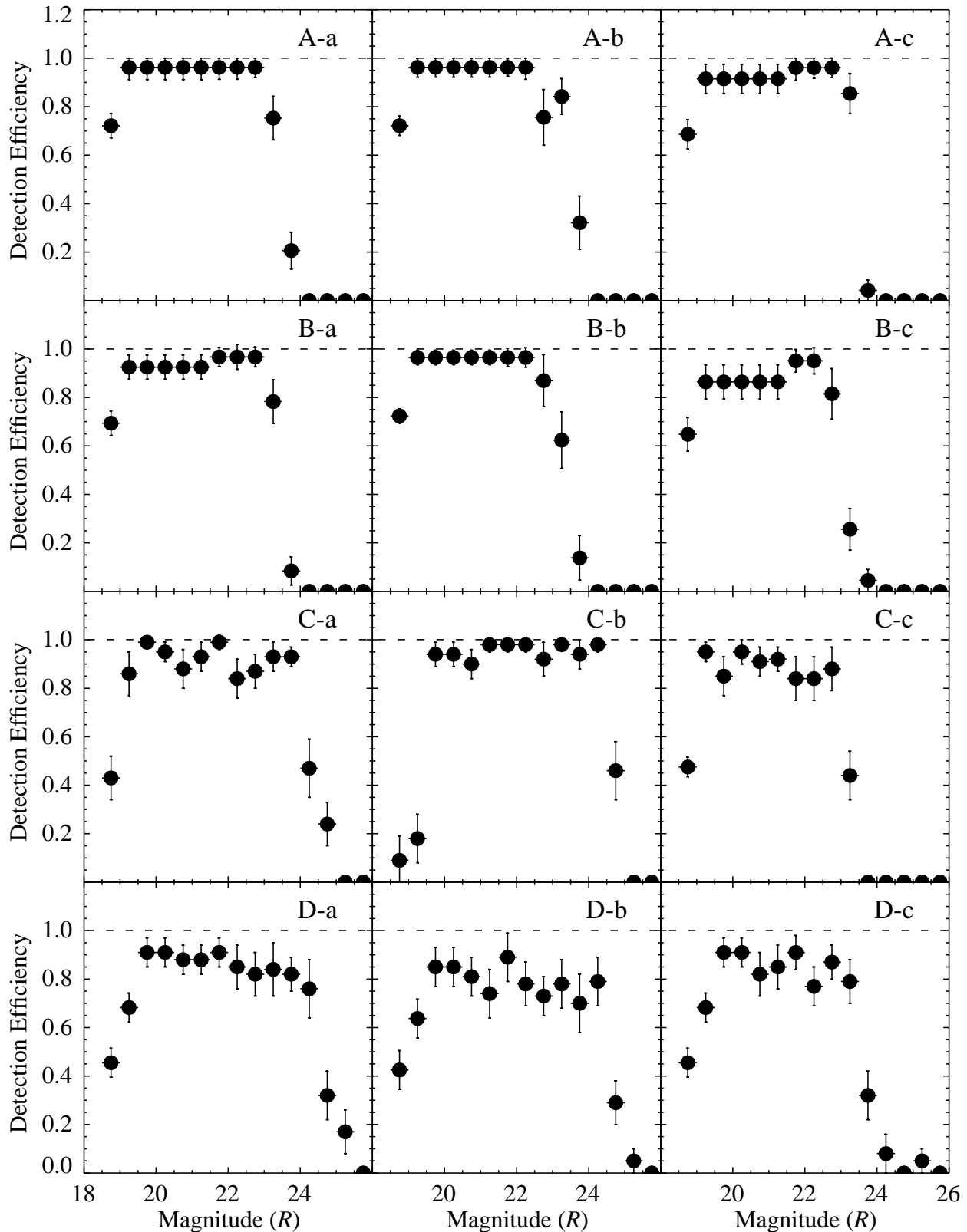


Fig. 1.— Detection efficiency versus  $R$  magnitude of the supernova for 12 representative examples among the 219  $2k \times 2k$  fields that were searched for SNe. Supernovae 1995as, 1996cj, 1997ai and 1997ep were respectively discovered on fields A-a, B-a, C-a and D-a.

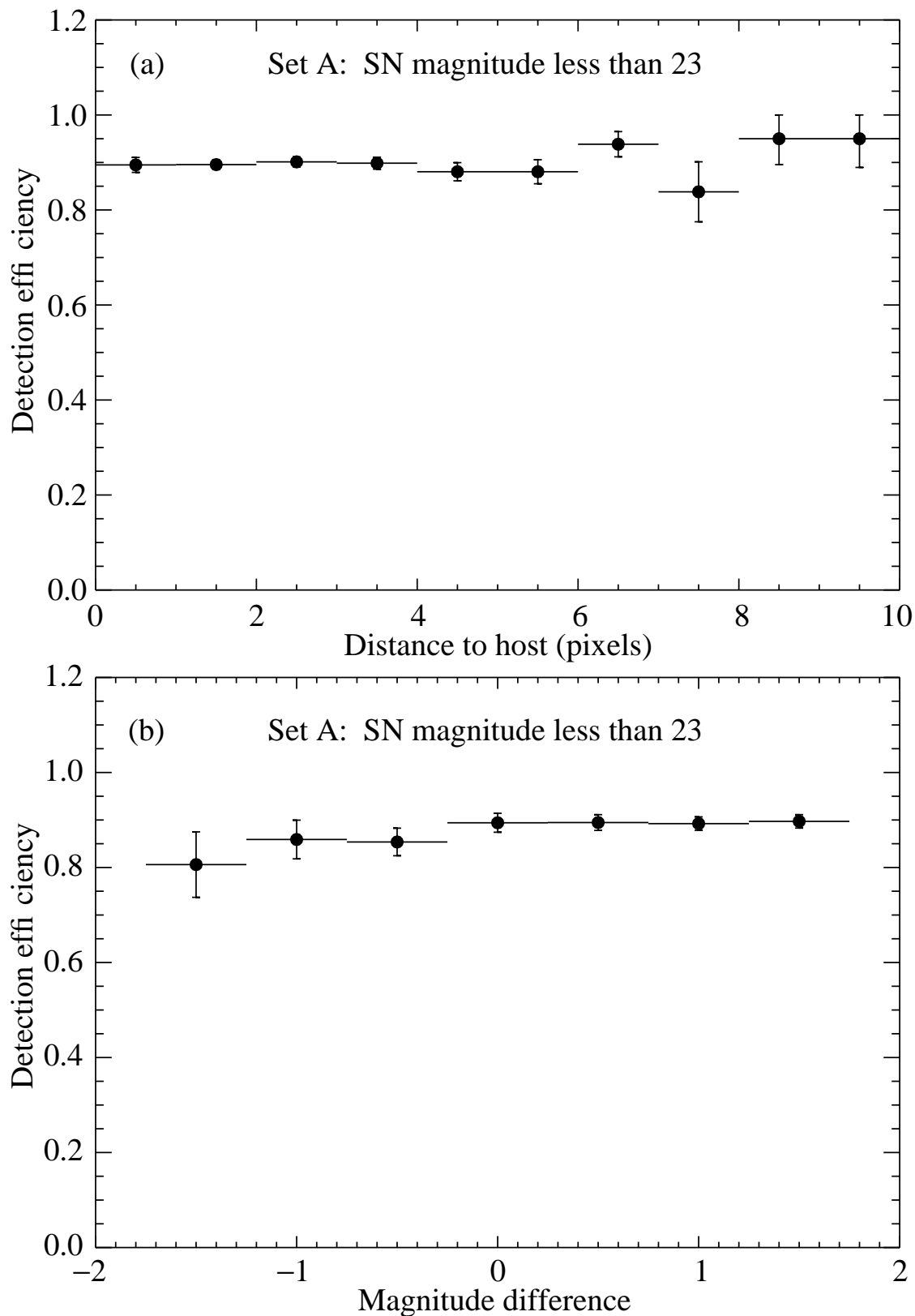


Fig. 2.— (a) detection efficiency versus projected distance to host galaxy; (b) detection efficiency as a function of magnitude difference between the host galaxy and the supernova (Host  $R$ -mag - SN  $R$ -mag). In both plots, an overall  $\sim 10\%$  inefficiency is present, due to the areal coverage lost by masking the region surrounding bright stars, independently of the

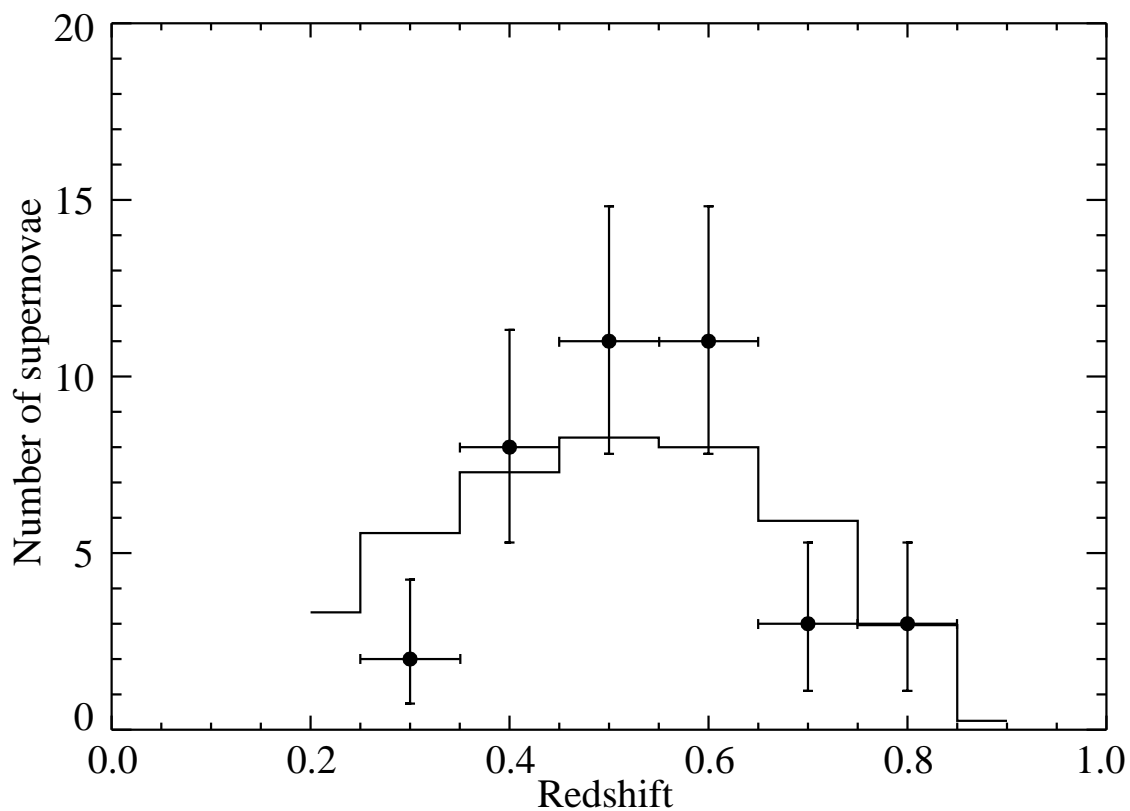


Fig. 3.— SN Ia Rate per unit comoving volume: Comparison of Monte-Carlo calculation (histogram) and data (points) for the number of observed SNe as a function of redshift. A value of  $1.53 h^3 10^{-4} \text{ Mpc}^{-3} \text{ yr}^{-1}$  is assumed for the rate. The prediction assumes no evolution for the rate per unit comoving volume computed with  $\Omega_M = 0.28$  and  $\Omega_\Lambda = 0.72$ .

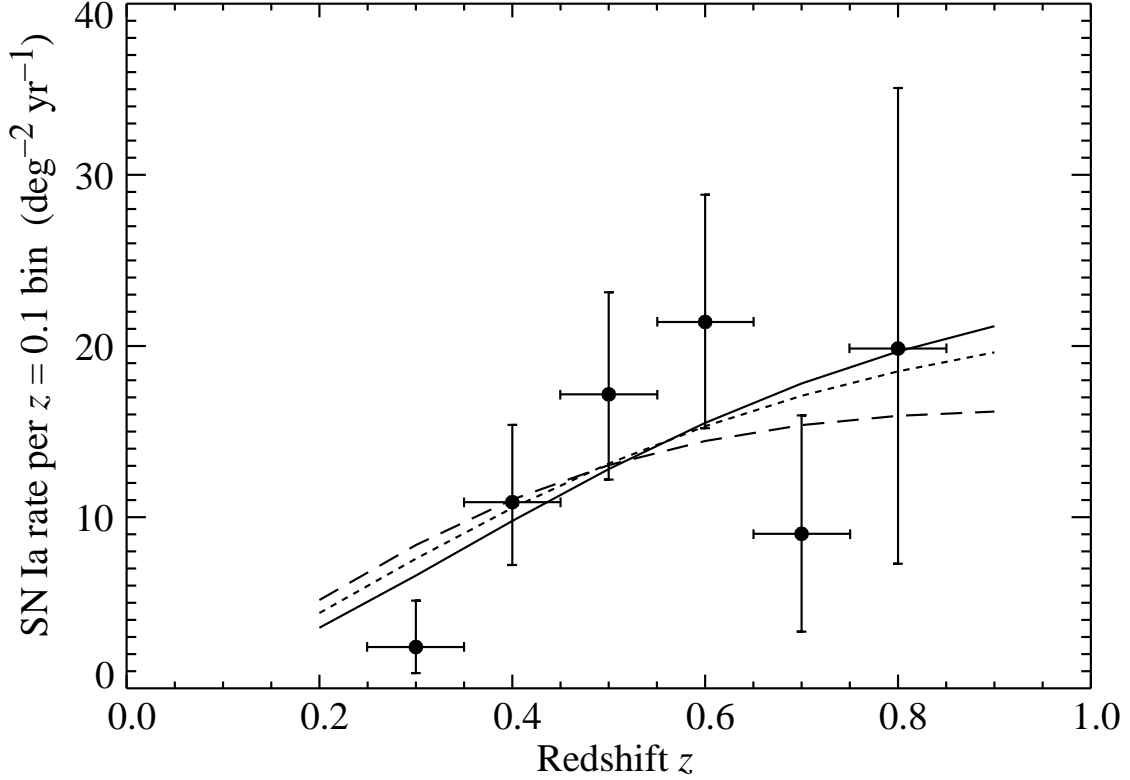


Fig. 4.— Number of SNe Ia per  $z=0.1$  redshift bin per square degree per year as a function of redshift. Over plotted (lines) are the predictions assuming that the number of SNe is proportional to the comoving volume and adjusted to best fit the observed number of SNe between  $z = 0.25$  and  $z = 0.85$ . The solid line is for a comoving volume given by a cosmological model with  $\Omega_M = 0.28$  and  $\Omega_\Lambda = 0.72$  while the dotted line is for  $\Omega_M = 0.3$  and  $\Omega_\Lambda = 0.0$  and the dashed line is for  $\Omega_M = 1.0$  and  $\Omega_\Lambda = 0.0$ .

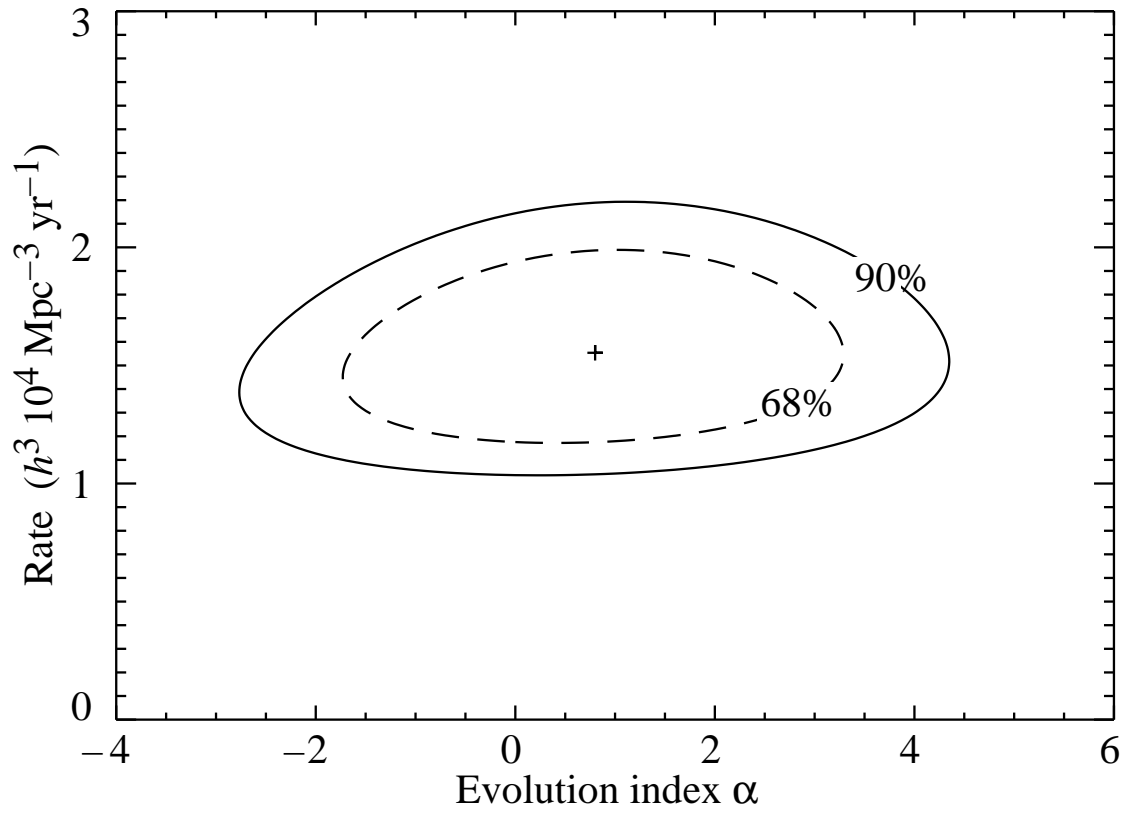


Fig. 5.— Two-parameter maximum likelihood fit of the distant SN Ia rate: 68.3% and 90% confidence regions for the rate per unit comoving volume versus the evolution index for a comoving volume corresponding to a flat universe with  $\Omega_M = 0.28$ .

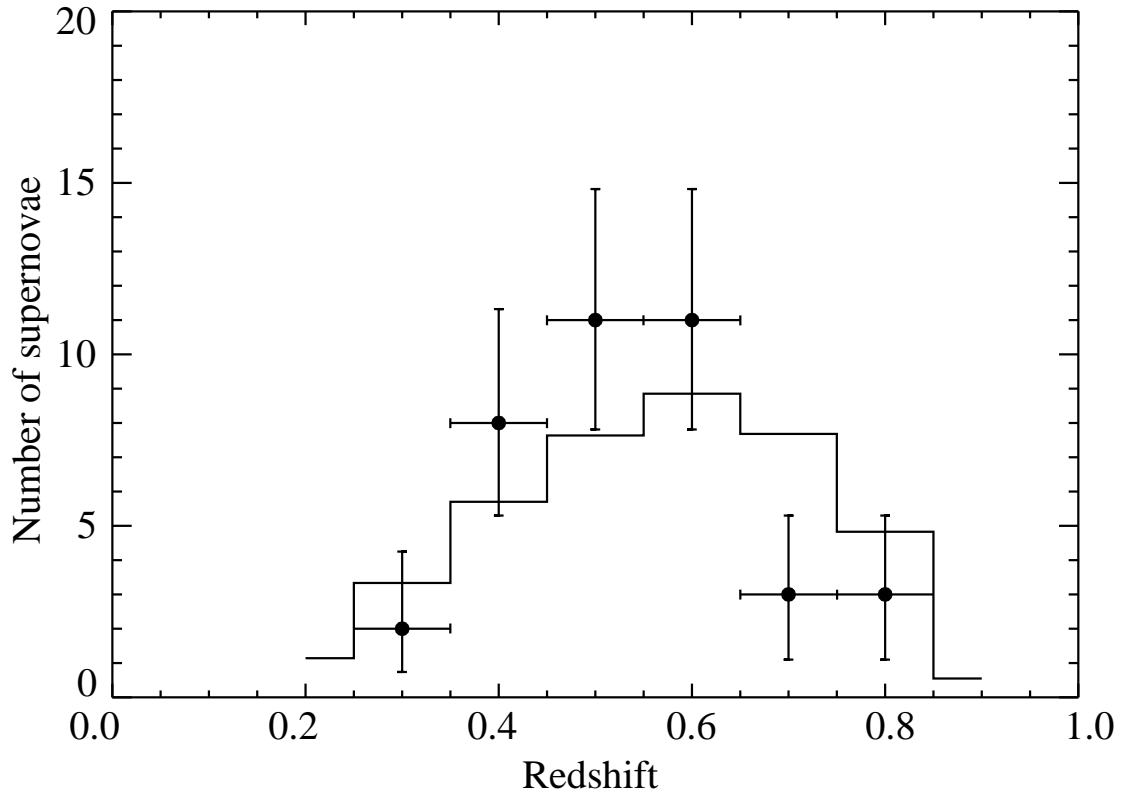


Fig. 6.— SN Ia Rate per unit luminosity: Comparison of Monte-Carlo calculation (histogram) and data (points) for the observed number of SNe as a function of redshift. The prediction assumes that the rate follows the galaxy luminosity evolution as a function of redshift. A value of  $0.58 h^2$  SNU is assumed for the rate, and  $\Omega_M = 0.28$  and  $\Omega_\Lambda = 0.72$  is used.

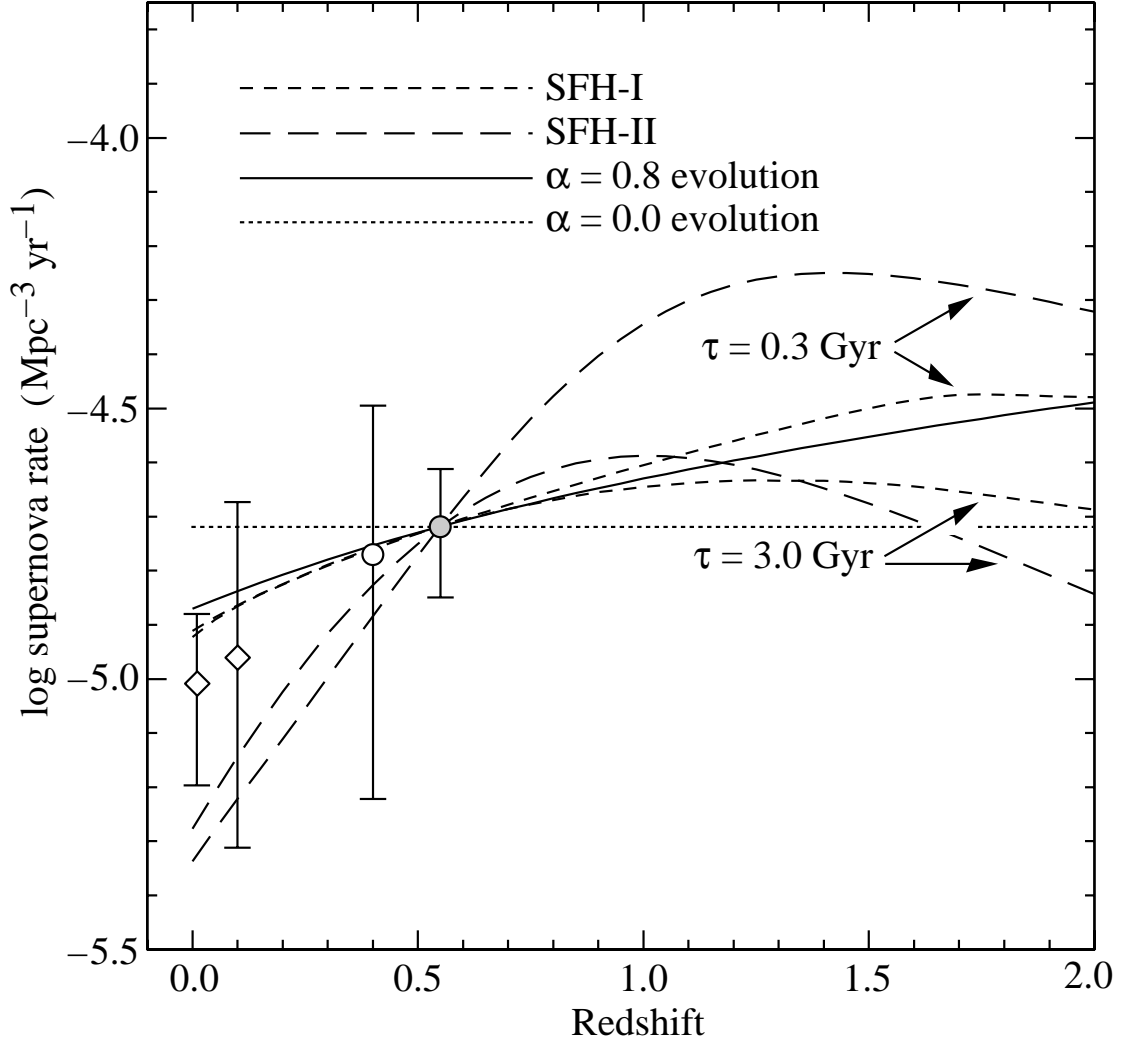


Fig. 7.— The SN Ia rate per comoving volume determined here (filled circle) compared with that of Paper I (open circle) and of Cappellaro et al. (1999) at  $z \sim 0.01$  and Hardin et al. (2000) at  $z \sim 0.1$  (open diamonds). For comparison, theoretical predictions for two star formation history scenarios and two delay times are shown, see text for details. Local SN Ia rates have been converted from SNU units. Also shown are an  $\alpha = 0.8$  evolution in the SNeIa rate (solid line) as well as the no evolution case (dotted line). The diagram is drawn for  $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}$  and a flat  $\Lambda$  dominated model with  $\Omega_M = 0.3$ .